

Measurement of β -delayed neutron emission probabilities for progenitors of the $A = 130$ r-process abundance peak

J. Liu,^{*1} V. H. Phong,^{*2,*3} A. Estrade,^{*4} G. Lorusso,^{*2,*5} F. Montes,^{*6} T. Davinson,^{*7} O. Hall,^{*7} K. Matsui,^{*2} N. Nepal,^{*4} S. Nishimura,^{*2} J. Agramunt,^{*8} D. S. Ahn,^{*2} A. Algora,^{*8} H. Baba,^{*2} N. T. Brewer,^{*9} C. Bruno,^{*7} R. Caballero-Folch,^{*10} F. Calvino,^{*11} I. Dillmann,^{*10} C. Domingo-Pardo,^{*8} S. Go,^{*12} C. J. Griffin,^{*7} R. Grzywacz,^{*9,*12} T. Isobe,^{*2} D. Kahl,^{*7} G. Kiss,^{*2} S. Kubono,^{*2} A. I. Morales,^{*8} B. C. Rasco,^{*9} K. P. Rykaczewski,^{*9} H. Sakurai,^{*2} Y. Shimizu,^{*2} T. Sumikama,^{*2} H. Suzuki,^{*2} H. Takeda,^{*2} J. L. Tain,^{*8} A. Tarifeño-Saldivia,^{*11} A. Tolosa-Delgado,^{*8} P. Woods,^{*7} and R. Yokoyama^{*12} for the BRIKEN collaboration

The first observation of a merger of two neutron stars,¹⁾ with both gravitational and electromagnetic wave signals, offers tantalizing opportunities to finally identify the astrophysical site of the r-process. The new observations will increase the demand for precise nuclear data necessary to reach a detailed understanding of the r-process mechanism. The r-process abundance peak around $A = 130$ is of particular interest because its shape and position is very sensitive to the the neutron-richness of the astrophysical environment, as its formation reflects the break-out of the reaction flow from the $N = 82$ classical waiting point isotopes. However, the effect on the final r-process abundance is obscured by a numbers of β -delayed neutron emitters along the decay path back to stability. In fact, the β -delayed neutron emission probabilities (P_n values) in the region south-east of ^{132}Sn have a most pronounced effect on the final r-process abundance, according to the recent sensitivity study in Ref. 2).

In June 2017, we have performed an experiment to study the decay properties of the β -delayed neutron emitters in the mass region $A = 130$ near the doubly magic nucleus ^{132}Sn . These neutron-rich isotopes were produced by the projectile fragmentation of a 345 MeV/nucleon ^{238}U beam on a Be target, before being purified and identified by the BigRIPS spectrometer. They were then transported through the Zero-degree spectrometer to reach the decay station located at the F11 focal plane. In the decay station, the active stopper array AIDA³⁾ was placed at a central position for the implantation of nuclei of interested, and it detected their subsequent β decay. The AIDA detector is a stack of six 8×8 cm² DSSDs with 128×128 pixels each. Neutrons emitted from the β decay of ions implanted in AIDA were detected by the BRIKEN neu-

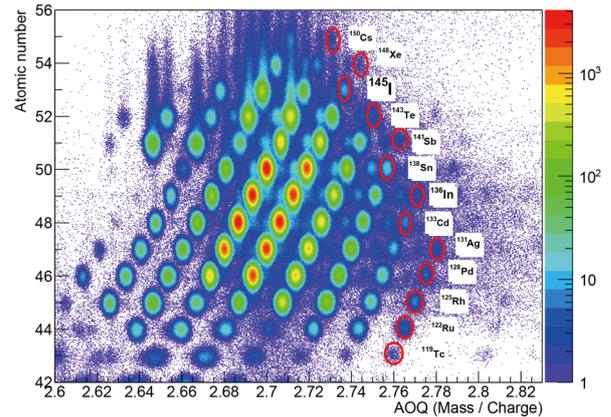


Fig. 1. Particle identification plot for isotopes transported to the decay station at F11, indicating the most neutron-rich isotopes for each element.

tron detector array⁴⁾ consisting of 140 gas-filled ^3He counters, which were inside a large moderation block made of high-density polyethylene. In addition, two clover-type high-purity Germanium detectors were employed to measure β -delayed and isomeric γ rays.

The particle identification plot combining data of the two settings of the experiment, centered at ^{130}Ag and ^{140}Xe , is shown in Fig. 1. The data analysis is ongoing. Preliminary results indicate that new or improved measurements of P_n values will be obtained for over 40 isotopes, and for 11 isotopes in the case of half-lives. These new measurements would make a significant contribution to the available experimental data for r-process models.

References

- 1) Abbott *et al.*, Phys. Rev. Lett. **119**, 161101 (2017).
- 2) M. R. Mumpower, R. Surman, G. C. McLaughlin, A. Aprahamian, Prog. Part. Nucl. Phys. **86**, 86 (2016).
- 3) C. J. Griffin *et al.*, Proc. XIII Nuclei in the Cosmos **1**, 97 (2014).
- 4) A. Tarifeño-Saldivia *et al.*, J. Instrum. **12**, 04006 (2017).

*1 Department of Physics, The University of Hong Kong
 *2 RIKEN Nishina Center
 *3 Faculty of Physics, VNU Hanoi University of Science
 *4 Department of Physics, Central Michigan University
 *5 National Physics Laboratory(NPL)
 *6 National Superconducting Cyclotron Laboratory
 *7 School of Physics and Astronomy, University of Edinburgh
 *8 Instituto de Fisica Corpuscular, Universidad de Valencia
 *9 Oak Ridge National Laboratory
 *10 TRIUMF
 *14 Universitat Politècnica de Catalunya
 *12 Department of Physics and Astronomy, University of Tennessee